

SXP 5.05 = IGR J00569-7226 : using X-rays to explore the structure of a Be star’s circumstellar disk

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Oct 31 2014

ABSTRACT

On MJD 56590-1 (2013 Oct 25-26) observations of the Magellanic Clouds by the *INTErnational Gamma-Ray Astrophysics Laboratory* (*INTEGRAL*) observatory discovered a previously-unreported bright, flaring X-ray source. This source was initially given the identification IGR J00569-7226. Subsequent multi-wavelength observations identified the system as new Be/X-ray binary system in the Small Magellanic Cloud. Follow-up X-ray observations by *Swift* and *XMM-Newton* revealed an X-ray pulse period of 5.05s and that the system underwent regular occultation/eclipse behaviour every 17d. This is the first reported eclipsing Be/X-ray binary system in the SMC, and only the second such system known to date. Furthermore, the nature of the occultation makes it possible to use the neutron star to “X-ray” the circumstellar disk, thereby, for the first time, revealing direct observational evidence for its size and clumpy structure. *Swift* timing measurements allowed for the binary solution to be calculated from the Doppler shifted X-ray pulsations. This solution suggests this is a low eccentricity binary relative to others measured in the SMC. Finally it is interesting to note that the mass determined from this dynamical method for the Be star ($\sim 13.0M_{\odot}$) is significantly different from that inferred from the spectroscopic classification of B0.2Ve ($\sim 16.0M_{\odot}$) - an effect that has been noted for some other high mass X-ray binary (HMXB) systems.

Key words: stars:neutron - X-rays:binaries

1 INTRODUCTION AND BACKGROUND

The Be/X-ray systems represent the largest sub-class of all HMXBs. A survey of the literature reveals that of the ~ 240 HMXBs known in our Galaxy and the Magellanic Clouds (Liu et al., 2005, 2006), $\geq 50\%$ fall within this class of binary. In fact, in recent years it has emerged that there is a substantial population of HMXBs in the Small Magellanic Cloud (SMC) comparable in number to the Galactic population. Though unlike the Galactic population, all except one of the SMC HMXBs are Be star systems. In these systems the orbit of the Be star and the compact object, presumably a neutron star, is generally wide and eccentric. X-ray outbursts are normally associated with the passage of the neutron star close to the

circumstellar disk (Okazaki & Negueruela 2001), and generally are classified as Types I or II (Stella, White & Rosner, 1986). The Type I outbursts occur periodically at the time of the periastron passage of the neutron star, whereas Type II outbursts are much more extensive and occur when the circumstellar material expands to fill most, or all of the orbit. General reviews of such HMXB systems may be found in Reig (2011), Negueruela (1998), Corbet et al. (2009) and Coe et al. (2000, 2009).

One of the aims of the *INTEGRAL* SMC survey is the ongoing study of the Be/X-ray binary population of the SMC, which can be used as a star formation tracer for ~ 50 Myr old populations (Antoniou et al. 2010). In this paper we present the analysis of X-ray and optical data from a newly discov-

ered X-ray source which we catalogue as IGR J00569-7226 (Coe et al., 2013b). This source exhibits most of the characteristics of Be/X-ray binary systems, revealing both a binary period of 17d (Coe et al., 2013c) and an X-ray pulse period of 5.05s (Coe et al., 2013d). Because of this pulse period we will refer to the system as SXP 5.05 throughout this paper (following the naming convention of Coe et al., 2005). Unusually, the neutron star in this system is clearly occulted by the Be star's circumstellar disk. In addition, there is evidence for a complete eclipse of the neutron star by the Be star - if so, this will be the first eclipsing Be/X-ray system in the SMC. The only previously known "eclipsing" Be/X-ray binary system is the recently discovered system, LXP 169, in the LMC (Maggi et al, 2013). But that system only exhibits transits by the neutron star across the face of the Be star and no actual eclipses by the Be star of the neutron star - presumably possible if the orbital eccentricity is large enough.

Shown in Figure 1 is an overall summary of the behaviour of this source at optical and X-ray wavelengths. Each data set is discussed separately in the following sections and possible models presented in Section 4.

2 X-RAY OBSERVATIONS

2.1 INTEGRAL

SXP 5.05 was detected throughout the 2013 SMC monitoring campaign with the *INTEGRAL* observatory (Winkler et al. 2003). The Imager on Board the *INTEGRAL* satellite (IBIS) (Ubertini et al. 2003) is optimised for an energy range of 15–200 keV and has a field of view of 29×29 degrees (Full Width Zero Intensity response), and so is well suited to observing large sky areas for point sources. The Joint European X-ray Monitor (JEM-X) telescope provides additional spectral coverage down to 3 keV over a smaller field of view. Observations were performed approximately every two *INTEGRAL* orbits, and typically consisted of 2 repetitions of the standard 5×5 dither pattern within each of those orbits. Overall the observations spanned the period from revolution 1347 (MJD 56590, 2013 Oct 24) to revolution 1373 (MJD 56668, 2014 Jan 11). Total IBIS exposure was 888ks, with an overall live-time of 620ks in ~ 450 pointings.

In order to provide rapid source discovery, near real-time data were processed using the *INTEGRAL* Offline Science Analysis (OSA, Goldwurm et al. 2003) v9 software and instrument characteristics files. IBIS (17–35 keV) and JEM-X (3–10 keV) maps were constructed for each revolution of data and searched for excesses. The IBIS energy band was chosen to maximise the detection significance of SMC X-1, and hence other SMC accreting X-ray pulsars, which have similar spectral shapes to SMC X-1 in this energy range. The consolidated data were re-processed using the latest OSA v10 software when they became available.

SXP 5.05 was first seen in revolution-level mosaics from *INTEGRAL* revolution 1347 (Oct 24th 2013), at a detection significance of 7.6σ in the IBIS mosaic and 9.5σ in the JEM-X mosaic. An initial position estimate was taken from the combined JEM-X1/JEM-X2 mosaic using the `j_src_locator` tool, and was (RA,dec)=(14.239,-72.431) with a position uncertainty $\sim 2'$.

In the following observations in revolution 1349 (Oct

30th 2013), SXP 5.05 was observed to have faded significantly in both IBIS and JEM-X, dropping to 3σ in the IBIS map (this is below the usual source detection limits) and 4.9σ in JEM-X. The changes in the IBIS and JEM-X broad-band fluxes between revolutions 1347 and 1349 were indicative of a change in spectrum of the source.

In revolution 1352 (9 Nov 2013), SXP 5.05 was seen to brighten again, reaching its highest detected flux levels in both 3–10 keV and 17–35 keV bands, both of which were $\sim 25\%$ higher than the discovery levels in rev 1347. Thereafter, the source flux was seen to reduce slowly in both IBIS and JEM-X energy bands. By the time of revolution 1373 (Jan 11 2014), the source was once again formally below the IBIS detection levels (at 4.3σ) but still detected in JEM-X (4.6σ). The overall evolution of the fluxes from IBIS and JEM-X (combined JEM-X1 and JEM-X2 maps) are shown in Figure 1.

Shown in Figure 2 are the images from the summed total data set of the *INTEGRAL* monitoring campaign. From these images four sources are clearly visible and one more lies in the Wing of the SMC:

(a) IGR J00569-7226 (=SXP 5.05) which is the subject of this paper.

(b) SMC X-1 which varied in brightness over the 2-3 months of observing sessions following the familiar super-orbital and binary modulation (Coe et al. 2013a).

(c) SXP 2.76 (= RX J0059.2-7138) is a well-established source originally identified by Hughes (1994) and Southwell & Charles (1996).

(d) IGR J01054-7253 (=SXP11.5) another well-established source seen in previous *INTEGRAL* surveys of the SMC (Coe et al. 2010).

(e) IGR J01217-7257 a bright previously-unknown source seen in just the last observation (11 January 2014) lying in the Wing of the SMC. Follow-up *Swift* observations located this source at a position of RA(J2000) = $01^h21^m41^s$, Dec(J2000) = $-72^\circ57'33''$, with an estimated uncertainty of $4''$ radius (Coe et al, 2014). This position was coincident with an OGLE IV optical object SMC732.03.3540 which revealed an 84d modulation. No X-ray pulsations were seen.

2.2 XMM-Newton

2.2.1 Observations

The *INTEGRAL* monitoring campaign came with three supporting *XMM-Newton* Target of Opportunity (ToO) observations. Data from the European Photon Imaging Cameras (EPIC) and the Reflection Grating Spectrometer (RGS) were processed using the *XMM-Newton* Science Analysis System v13.0.0 (SAS) along with software packages from FTOOLS v6.15.1. Table 1 summarises the details of the EPIC MOS (Turner et al. 2001) and pn (Strüder et al. 2001) observations.

The first observation (referred to as Obs1 hereafter) was performed MJD 56597 (2013 November 1), 7 days after the initial *INTEGRAL* detection. The *INTEGRAL* JEM-X count rate, 1.6 counts s^{-1} , implied an EPIC-pn 0.2–10.0 keV count rate of ~ 19 counts s^{-1} (assuming a simple absorbed power law model with $n_H = 6 \times 10^{20} \text{ cm}^{-2}$ and $\Gamma = 1.0$): above the 2 counts s^{-1} pile up threshold for the detector in full frame

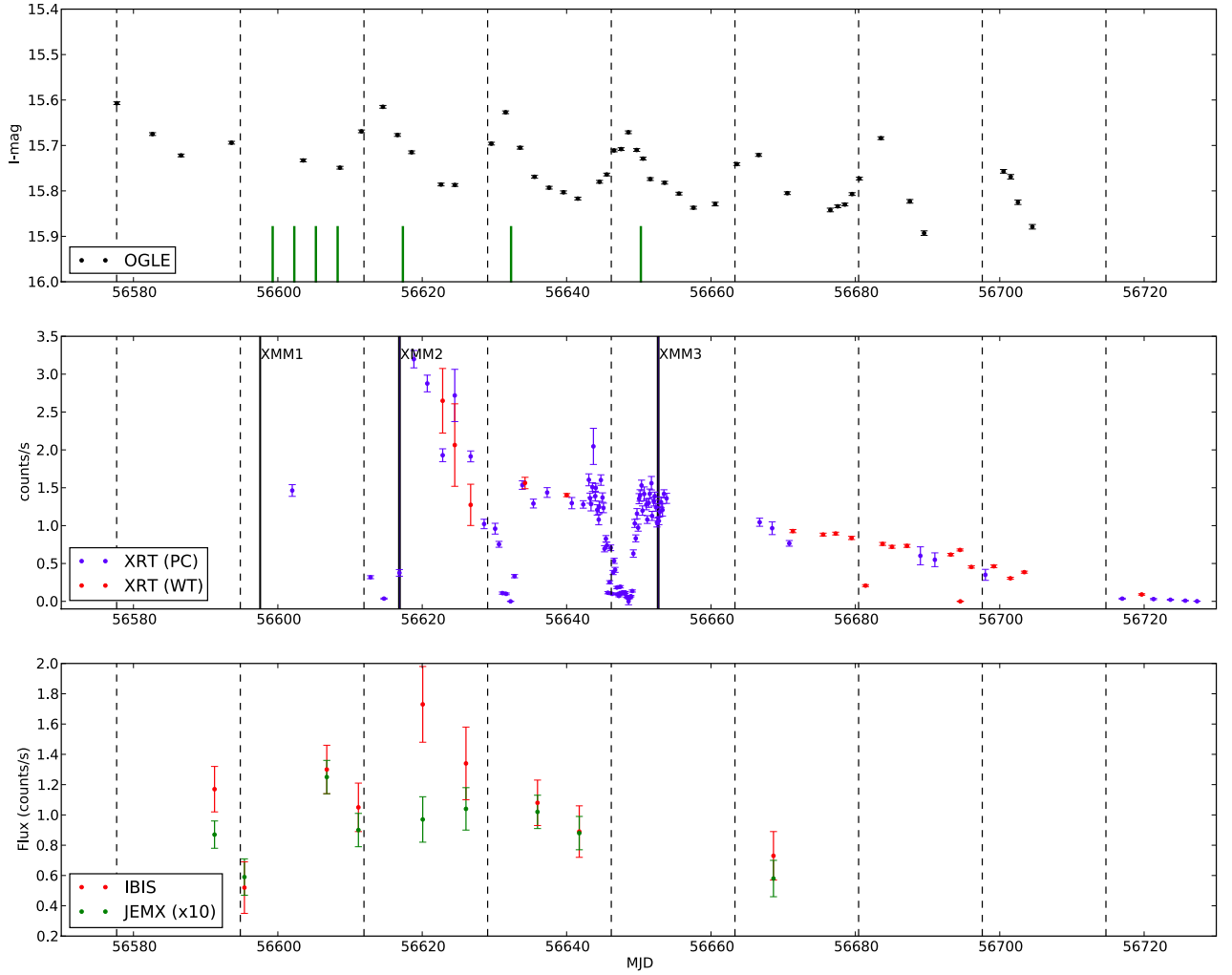


Figure 1. Summary of all observations of SXP 5.05. The dashed vertical lines indicate the position of periastron of the neutron star based upon the ephemeris presented in Table 5. Top panel: OGLE I-band photometry with the timings of the SALT spectroscopic measurements indicated as green vertical bars. Middle panel: *Swift* X-ray flux measurements with the positions of the detailed XMM-Newton measurements indicated by blue vertical lines. Bottom panel: *INTEGRAL* hard X-ray measurements from both the IBIS and JEM-X instruments.

Table 1. XMM-Newton EPIC observations of SXP 5.05.

ObsID	Camera	Filter	Read Out Mode	Observation Date (MJD)	Start(UT)	End(UT)	Exp. ^(a) (ks)	Count Rate ^(b) counts s ⁻¹
0700580101	MOS1	Thin	Full Frame	56597	10:39	14:24	13.5	0.297 ± 0.004
	MOS2	Thin	Full Frame		10:40	14:24	13.5	0.286 ± 0.004
	pn	Thin	Small Window		11:53	14:27	9.2	0.96 ± 0.01
0700580401	MOS1	Thin	Full Frame	56616	17:15	22:41	19.6	1.95 ± 0.01
	MOS2	Thin	Full Frame		17:15	22:41	19.6	1.94 ± 0.01
	pn	Thin	Full Frame		17:37	22:49	18.7	7.02 ± 0.02
0700580601	MOS1	Thin	Full Frame	56652	13:49	19:00	18.6	2.71 ± 0.01
	MOS2	Thin	Full Frame		13:50	19:00	18.6	2.69 ± 0.01
	pn	Thin	Small Window		13:55	19:05	18.5	11.66 ± 0.03

Notes. ^(a) Before any filtering has been applied. ^(b) 0.2-10.0 keV. From source detection, i.e. after filtering has been applied.

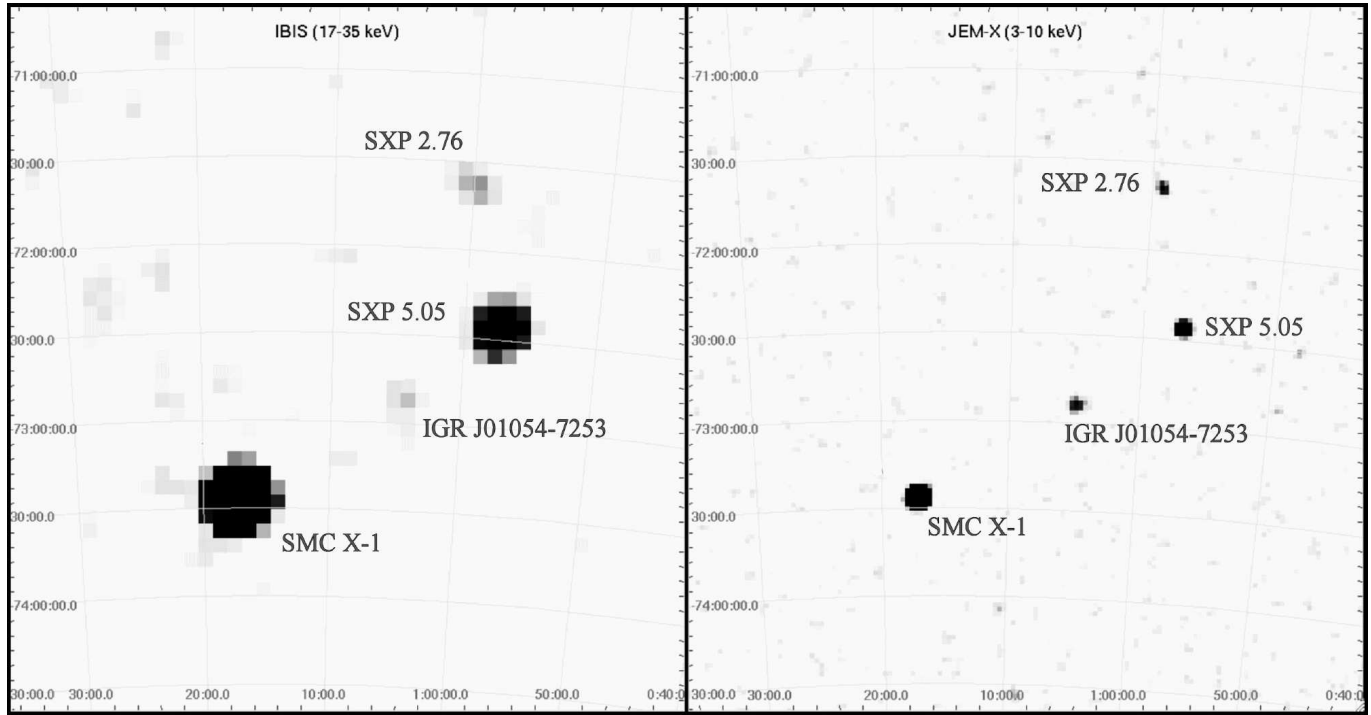


Figure 2. Summed *INTEGRAL* images from November 2013 - January 2014 observations covering 4×4 degrees in size. Left: IBIS 17-35 keV significance image. Right: combined JEM-X1 + JEM-X2 3-10 keV significance image.

mode¹. To avoid pile up, the EPIC-pn was used in small window mode which offers superior timing resolution but with a reduction in the field of view ($4.4' \times 4.3'$ as opposed to $\sim 30'$ diameter) and ontime of the telescope. The field of view of the EPIC-pn detector is similar to the 1σ error circle of the *INTEGRAL* source position (see Section 2.1). This presented a serious problem for source localisation. However, we were made aware of a new transient, with position consistent with the *INTEGRAL* error circle, detected in the *Chandra* large survey of the SMC (Haberl, priv comm.). The decision was made to use this position as the telescope pointing.

The count rate of the source in Obs1 was considerably less than expected (see Table 1). As such, the decision was made to leave the EPIC-pn detector in the default full frame readout mode for the second observation (Obs2), taken MJD=56616 (2013 November 20). The final observation (Obs3) of SXP 5.05 by *XMM-Newton* was made MJD=56652 (2013 December 26). Again, based on the source activity in the period since the Obs2, the EPIC-pn detector was returned to small window mode. Figure 3 shows false colour images of the three EPIC observations which reveal the detection of 4 other Be/X-ray systems in the same field as SXP 5.05. One of these objects, SXP 7.92, is recent *Chandra* discovery (Israel et al, 2013) and these *XMM-Newton* data will be the subject of a more detailed paper (Israel et al, 2014).

2.2.2 Data Reduction

The EPIC-MOS and pn observational data files were processed with *emproc* and *epproc* respectively. The data were

screened for periods of high background activity by examining the >10 keV count rate. No filtering was required for the first and third observation. The first 5.9 ks of Obs2 were affected by high background. These intervals were removed from our analysis.

Cleaned images were then created in 5 energy bands: 0.2-1.0 keV, 1.0-2.0 keV, 2.0-4.5 keV, 4.5-10.0 keV and 0.2-10.0 keV for each observation. A box sliding detection was performed simultaneously on all 12 images (4 bands \times 3 detectors) twice (the first with a locally estimated background the second using the background map) with the task *eboxdetect*, followed by the maximum likelihood fitting using the task *emldetect*. This process resulted in a list of sources, including their positions, errors and background subtracted counts, for each of the three observations.

As well as searching for sources in each observation individually, we also performed the source detection algorithm on all three observations simultaneously (i.e. 36 images) using the tasks *emosaic_prep* and *emosaicproc*.

The average count rate of Obs2 exceeds the threshold for pile up for the read out mode. The observation was checked for pile up using the SAS task *epatplot*. Despite the high count rate no evidence for pile up was found.

Light curves and spectra of SXP 5.05 were extracted from the EPIC-pn detector with the same energy ranges as described above. Despite the variations in the source count rate, a extraction radius of $30''$ was used in all observations to avoid any contamination from the neighbouring source. "Single" and "double" pixel events were selected with quality flag (FLAG=0), i.e. all bad pixels and columns were disregarded. Photon arrival times were converted to barycentric dynamical time, centred at the solar system barycenter, using the SAS task *barycen*. For the first and third observation, the

¹ <http://xmm.esac.esa.int/external/xmm.user.support/documentation/uhb/epicmode.html>

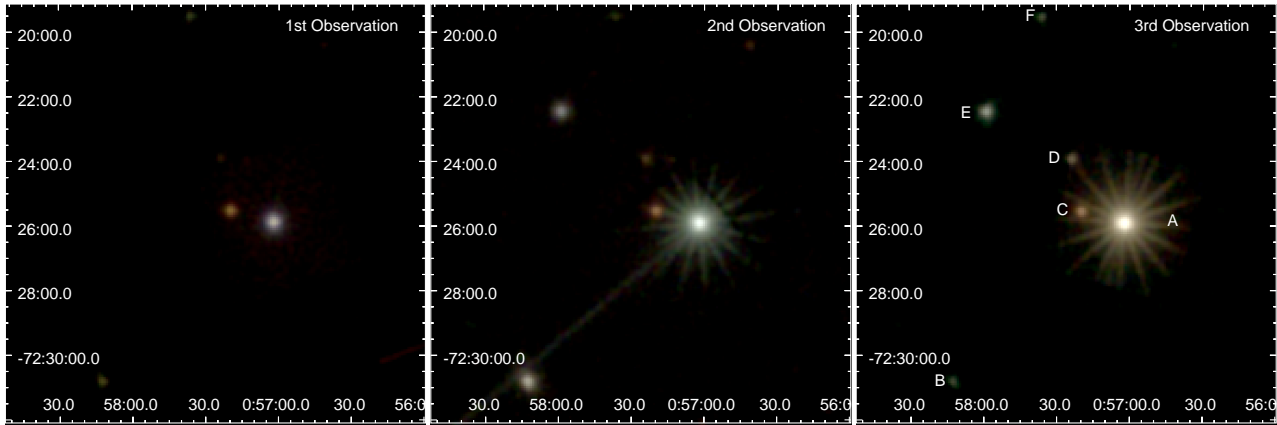


Figure 3. False colour mosaic images of the three EPIC detectors. Red corresponds to 0.2-2.0 keV, green to 2.0-4.5 keV and blue to 4.5-10.0 keV. All images have had a Gaussian smoothing applied, with a width of 2 pixels, and have the same colour scale for comparison. Sources identified in the field of view are: **A.** SXP 5.05 (the subject of this paper), **B.** SXP 293, **C.** A known quasar with $z=0.15$, identified in both the *XMM-Newton* and *Chandra* point source catalogues, **D.** HMXB candidate XMMU J005724.0-722357, **E.** SXP 7.92, a recently identified Be/X-ray binary system (Israel et al, 2013), **F.** SXP 565.

Table 2. Position of SXP 5.05 as determined by *XMM-Newton*

Obs	RA (J2000)	dec (J2000)	σ_{stat} arcsec
Obs1	$00^h 57^m 2.24^s$	$-72^\circ 25' 53.7''$	0.2
Obs2	$00^h 57^m 2.20^s$	$-72^\circ 25' 54.35''$	0.02
Obs3	$00^h 57^m 2.32^s$	$-72^\circ 25' 54.99''$	0.01
Combined	$00^h 57^m 2.30^s$	$-72^\circ 25' 54.75''$	0.02

background region fell on the same CCD chip and had an extraction radius of $30''$ (identical to that of the source region). For Obs2, the background region fell on a neighbouring CCD in the pn detector. The extraction radius was $100''$.

The count rate of the source was high enough in the second and third observation to allow RGS spectra to be extracted. Despite the flare identified at the start of Obs2 in the EPIC instruments, no filtering was required for either instrument for either observation.

2.2.3 X-ray Position

The source detection described in Section 2.2.1 was performed on all three EPIC cameras for each of the *XMM-Newton* observations, both individually and simultaneously. Table 2 lists the position of SXP 5.05 determined in each observation as well as the final position determined from fitting all 45 images simultaneously, along with the 1σ statistical error, σ_{stat} . The 1σ systematic uncertainty was assumed to be 1 arcsec in accordance with the findings of the *XMM-Newton* Serendipitous Source catalogue (Watson et al. 2009). This is the dominant component of the error on the positions. These positions are consistent with that reported by Kennea (2013). Figure 4 shows a *I*-band image with the location of the *Swift* and *XMM-Newton* positions with radii equal to the 1σ errors.

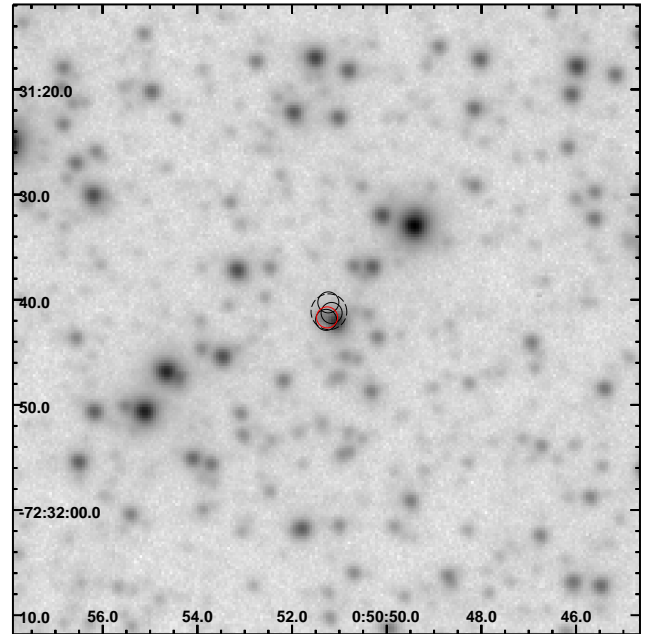


Figure 4. A 1×1 arcmin *I*-band image of SXP 5.05 taken from the OGLE IV data. The smaller, solid black circles represent the 1σ error circles of the three individual *XMM-Newton* ToOs. The red solid circle is the source position as determined by simultaneous source detection on all three observations. The broken circle is the *Swift* 90% confidence radius.

2.2.4 Timing Analysis

Figure 5 shows the EPIC-pn 0.2-10.0 keV light curves of all three observations alongside each other. Each observation is characterised by its own distinctive behaviour (see Table 3).

Figure 6 shows the Lomb-Scargle periodograms of the EPIC-pn 0.2-10.0 keV 0.1 s binned light curves and the normalised pulse profiles of each observation. Unusually for a Be/X-ray, the light curve of Obs1 shows no evidence for pul-

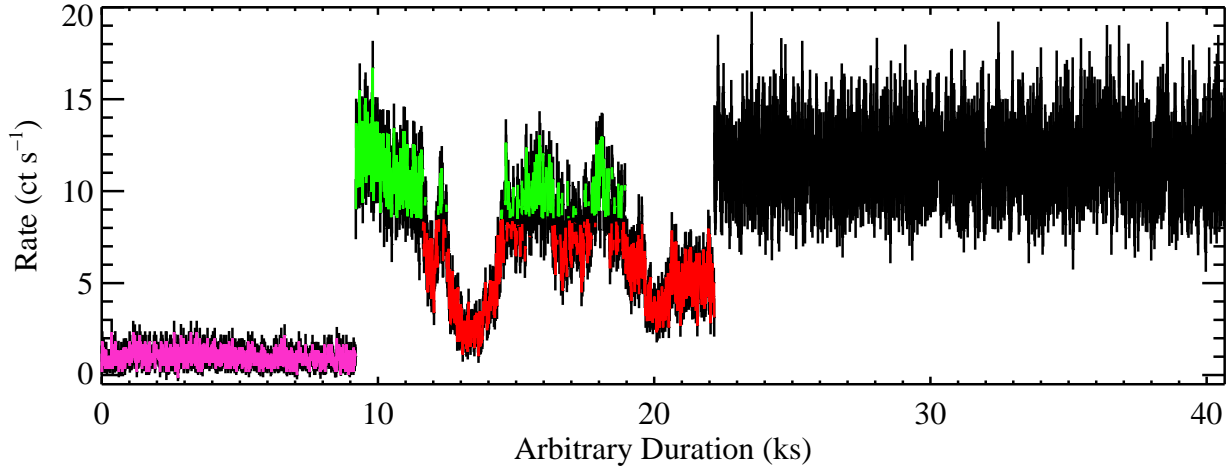


Figure 5. The three separate *XMM-Newton* EPIC-pn 0.2-10.0 keV light curves concatenated together. Obs1 is magenta, high count rate intervals of Obs2 are in green, low count rate intervals of Obs2 in red and Obs3, black. The colours correspond to the spectra shown in Section 2.2.5

Table 3. Summary of the timing analysis of all 3 *XMM-Newton* observations. The fractional root-mean-squared value has even calculated according to Equation (10) of Vaughan et al. (2003)

Obs	Count Rate ct s ⁻¹	Fractional rms	Pulse Period (s)	Pulsed Fraction
Obs1	0.95 ± 0.01	0.11	-	<0.35 ^(a)
Obs2	7.36 ± 0.03	0.39	5.05066 ± 0.00003	0.26 ± 0.03
Obs3	11.60 ± 0.03	0.09	5.04843 ± 0.00001	0.56 ± 0.06

Notes. ^(a) 3 σ upper limit.

sations. Pulsations are clearly detected in Obs2 and Obs3 at 5.051 s and 5.048 s with Lomb-Scargle power 833.4 and 5644.8 respectively.

Monte Carlo simulations were used to determine the errors on the pulse periods: 1000 new light curves were created for both Obs2 and Obs3 by varying each point in the original light curves within their 1σ errors using a Gaussian random number generator. Each of these light curves were searched for a period between 0.1 and 10 s in the same manner as the original light curves (i.e. Lomb-Scargle analysis). The histograms of the resulting periods are well fit with Gaussian functions with mean 5.05066 and 5.04843 and standard deviations 3×10^{-5} and 1×10^{-5} . Thus we determine the pulse period of SXP 5.05 to be 5.05066 ± 0.00003 s, as of MJD=56616, and 5.04843 ± 0.00001 s, as of MJD=56652.

The pulsed fractions of the light curves were calculated by integrating over the pulse profile and taking the ratio of the pulsed component of the profile (i.e. the area of the pulse profile above the profile minimum) and the total count rate. The light curves of Obs2 and Obs3 were split into four energy bands (0.2-1.0 keV, 1.0-2.0 keV, 2.0-4.5 keV and 4.5-10.0 keV) to investigate if there was/is any energy dependence in the shape of the pulse profile or the pulsed fraction. Within each observation, the normalised pulse profiles (i.e. the pulsed profiles divided by the average count rate in the energy bands) above 1.0 keV are consistent within the 1σ

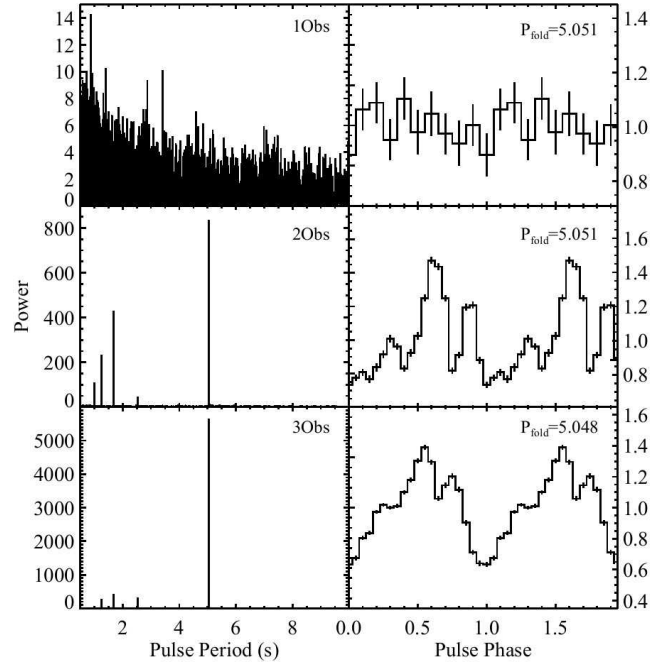


Figure 6. Lomb-Scargle periodograms and normalised pulse profiles of each *XMM-Newton* EPIC-pn 0.2-10.0 keV observation in descending order. The profiles have been arbitrarily aligned in phase. Obs1 shows no evidence for any pulsations and so the light curve has been folded on the period detected in Obs2.

errors. The 0.2-1.0 keV pulse profiles are consistent with the > 1.0 keV pulse profiles within the 2σ errors.

Both Obs2 and Obs3 were also investigated to see if there was any evidence for temporal variations in the hardness ratio (and hence the spectra). The hardness ratio is defined as:

$$\frac{C_{hard} - C_{soft}}{C_{hard} + C_{soft}} \quad (1)$$

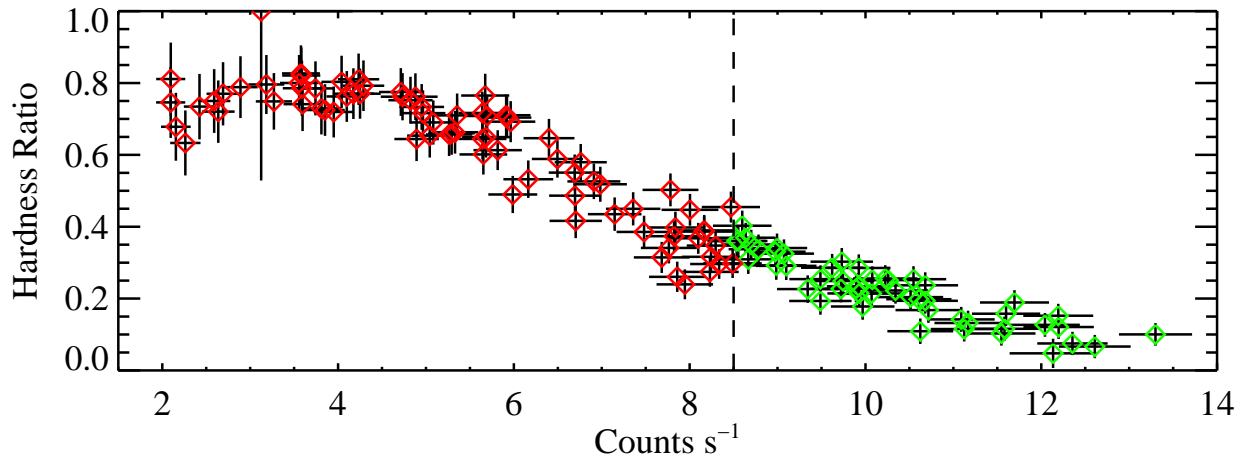


Figure 7. Variation of hardness ratio with count rate for Obs2. The broken line marks the count rate at which we split the data into hard and soft intervals

where C_{hard} and C_{soft} are the count rates in the 2.0-10.0 keV and 0.2-2.0 keV bands respectively. The hardness ratio can take any value between -1.0 (i.e. all source counts in the 0.2-2.0 keV energy range) and 1.0 (all counts in the 2.0-10.0 keV energy range). No significant variations were found on the timescale of the pulse period, however, Obs2 shows a clear anti-correlation between the total count rate of the source and hardness (see Figure 7): The source becomes appreciably harder at lower count rates. The light curve was split into two intervals, a soft/high count rate interval (indicated in green in Figures 5 and 7) a hard/low count rate interval for spectral analysis (see Section 2.2.5).

2.2.5 Spectral Analysis

To investigate the cause for the spectral changes between the three *XMM-Newton* observations and during the second observation, as indicated by the hardness ratio variations, we simultaneously fitted the EPIC-pn spectra with the same model, only allowing certain model parameters to vary. We included the total spectra from Obs1 and Obs3 and the two spectra from Obs2 (see Section 2.2.4), defined from hard/low count rate and soft/high count rate intervals. The spectra from Obs2 and Obs3 are very similar in shape and intensity at energies above 6 keV, indicating that absorption is the main parameter responsible for the spectral changes at lower energies.

To fit the spectra we started with a model consisting of a power-law and a soft component. The soft component dominates the spectra in the energy band of about 0.5 to 2 keV, depending on the amount of absorption seen in the power-law component. We tested a black-body model for this component, but found unrealistically high values for the size of the emission region to be compatible with an origin at the neutron star surface. Therefore, we used a multi-temperature accretion disk model (*diskbb* in XSPEC) which provides a similar fit quality. At very low energies (below 0.5 keV) another, even softer component is visible in the spectra. In eclipsing super-giant high mass X-ray binaries such soft emission is

visible due to scattering and re-processing of the X-ray photons originating from the vicinity of the neutron star in the strong stellar wind, when the direct X-ray emission along the line of sight is strongly suppressed by high absorption. This component was forced to be at an intensity of 2.5% of the direct emission components (including both power-law and disk blackbody), however, attenuated by lower absorption than the direct emission. To better constrain the parameters of the *diskbb* component (and to verify that it also represents higher-resolution spectra) we included the RGS spectra from Obs3 in the fit (all model parameters were linked with those for the corresponding EPIC-pn spectrum). No constant factor to allow for calibration uncertainties between the instruments was necessary to fit the RGS spectra. For each of the three emission components we applied absorption with free column densities in the fit with elemental abundances of 0.2 solar (for elements heavier than He) together with an overall absorption (fixed at $6 \times 10^{20} \text{ cm}^{-2}$ with solar abundances following Wilms et al. (2000) accounting for the Galactic foreground absorption. The model with its best fit parameters is summarised in Table 4. The fit results in a reduced χ^2_ν of 1.08 for 4461 degrees of freedom. The spectra together with best-fit model are shown in Fig. 8.

The simultaneous spectral modelling of the *XMM-Newton* observations demonstrates that the variability seen in the EPIC-pn light curve of Obs2 can solely be explained by absorption effects. The spectra obtained from Obs3 and the soft interval of Obs2 are very similar with comparable absorption column densities for the accretion disk and the power-law components. In contrast, the spectrum extracted from hard interval of Obs2 is strongly attenuated by a column density in excess of 10^{23} cm^{-2} for the direct power-law component.

Correcting the spectra for their absorption results in similar 0.2-10.0 keV luminosities between 4.3 and $5.2 \times 10^{37} \text{ erg s}^{-1}$ (assuming a distance of 60 kpc) for the spectra from Obs3, Obs2 hard and Obs2 soft, respectively. The RGS spectra reveal no significant emission lines which would be expected from optically thin gas surrounding the neutron star and re-

Table 4. Results of the simultaneous fit to the four EPIC-pn spectra described in the text, Obs1, Obs2 soft interval, Obs2 hard interval and Obs3. All errors are evaluated at the 90% confidence level. Fluxes and luminosities are given for the 0.2–10 keV band (distance 60 kpc).

Obs	$N_{H,db}^{(a)}$ 10^{22} cm^{-2}	$kT_{db}^{(a)}$ keV	$K_{db}^{(a)}$	$N_{H,pl}^{(b)}$ 10^{22} cm^{-2}	$\Gamma^{(b)}$	$K_{pl}^{(b)}$ $\times 10^{-3}$	$N_{H,sc}^{(c)}$ 10^{21} cm^{-2}	$K_{sc}^{(c)}$ %	F_x $\text{erg cm}^{-2} \text{ s}^{-1}$	L_x erg s^{-1}
Obs3	0.63 ± 0.03	0.42 ± 0.03	$23.9^{+8.1}_{-5.8}$	5.7 ± 0.5	1.53 ± 0.02	9.9 ± 0.3	$1.7^{+0.7}_{-0.4}$	2.5	6.0×10^{-11}	4.3×10^{37}
Obs2 soft	0.74 ± 0.04	-	-	9.9 ± 0.8	-	12.3 ± 0.4	0.1 ± 0.1	2.5	6.7×10^{-11}	5.2×10^{37}
Obs2 hard	5.8 ± 0.3	-	-	23.7 ± 1.6	-	10.0 ± 0.4	1.4 ± 0.2	2.5	4.4×10^{-11}	4.4×10^{37}
Obs1	$28.4^{+3.7}_{-3.1}$	-	-	85^{+14}_{-12}	-	1.7 ± 0.2	3.9 ± 0.4	2.5 ^d	6.8×10^{-12}	1.3×10^{37}

Notes. Spectral model in XSPEC convention: $phabs<1>(vphabs<2>*diskbb<3> + vphabs<4>*powerlaw<5> + vphabs<6>*(powerlaw<7> + diskbb<8>+gaussian<9>))$ (a) Direct disk blackbody component. Normalisation K_{db} in $(R_{in} / D_{10})^2 \times \cos(\theta)$ with inner disk radius R_{in} in km, distance D_{10} in units of 10 kpc and θ the angle of the disk ($\theta = 0$ is face-on). (b) Direct power-law component. Normalisation K_{pl} in [photons $\text{keV}^{-1} \text{ cm}^{-2} \text{ s}^{-1}$ at 1 keV]. (c) Scattered component with fixed normalisation relative to the direct component (which includes *powerlaw* and *diskbb*). ^d The scattered component for Obs1 was fixed relative to the direct components of Obs3.

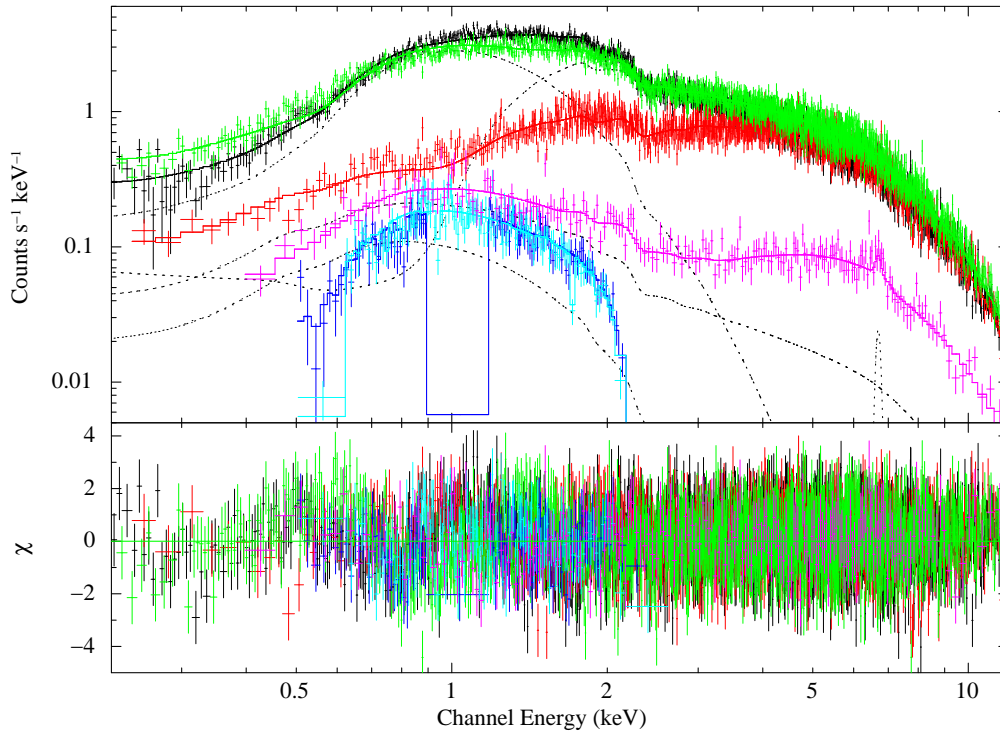


Figure 8. The X-ray spectra of SXP 5.05 as observed during Obs1 (magenta), Obs2 (hard/low count rate intervals in red, soft/high count rate intervals in green) and Obs3 (black) as observed by EPIC-pn. The RGS spectra of Obs3 are shown in blue and light-blue. The upper panel presents the spectra and best fit model (as solid line) while the residuals are plotted in the lower panel. The individual model components for Obs3 are indicated by dashed lines.

processing the hard X-ray emission. Hickox et al. (2004) argue that for brighter systems a soft excess in the X-ray spectrum is most likely due to radiation from the inner part of the accretion disk. The fact that the absorption column densities we derive from our spectral modelling is lower for the disk component than for the power-law component suggests the soft emission originates further from the neutron star, consistent with an optically thick accretion disk. From the normalisation of the *diskbb* model we derive an inner disk radius of 29.3 km, assuming an inclination of 90 degrees (i.e. accretion disk and orbital plane aligned and viewed close to edge-on).

The spectrum of Obs1 is much harder in comparison to the expected scattered component seen from the other two observations. A significant bump around 5 keV indicates a

relatively strong power-law component with very high absorption, even higher than that in the hard interval of Obs2. This is confirmed by the combined analysis of the EPIC-pn spectra. Overall, the spectrum of Obs1 is well reproduced by the same model, including the direct *diskbb* and *powerlaw* components, with extreme column densities (see Table 4), in addition to the scattered component (assumed to be the same as that seen in the spectrum of Obs3). The presence of direct emission is inconsistent with a full eclipse of the X-ray source by the Be star, but the extremely high absorption measured from the direct emission components suggests that the line of sight passes the Be star near its rim through the very dense innermost part of the Be disk. At high energies (above 6 keV) the power-law component is not affected by absorption,

but is subject to Thomson scattering which reduces the direct emission (this process is not included in the *phabs* model of XSPEC) significantly at column densities approaching 10^{24} cm^{-2} . The observed flux between 7.0 and 10.0 keV for Obs1 is 16% of that seen during Obs3. If the intrinsic source flux is similar during the two observations then an electron column density of $2.7 \times 10^{24} \text{ cm}^{-2}$ is required to explain the flux reduction.

Another spectral feature expected from absorption and reprocessing in the gas surrounding the X-ray source is Fe-K α fluorescence. A weak line is seen in the EPIC-pn spectrum of Obs1. The line energy was determined from the fit to $6.69 \pm 0.06 \text{ keV}$, consistent with an origin from highly ionised iron. The equivalent width for Obs1 is 78 eV, while for the other observations it is formally reduced to 10 eV by the strong underlying continuum. In the fit we assume a constant line strength, although no line is required for the spectra of Obs2 and Obs3 as the continuum completely dominates at energies around 6.7 keV. The equivalent width for Obs1 is $78 \pm 40 \text{ eV}$ (90% confidence), while for the other observations it is formally reduced to $12 \pm 7 \text{ eV}$ by the strong underlying continuum.

2.3 *Swift*

2.3.1 *Swift* archival observations

Including all data where SXP 5.05 was inside the *X-ray Telescope* (XRT) (Burrows et al., 2005) field of view (all observations were pointed within $8'$), *Swift* has serendipitously observed a region including SXP 5.05 on five occasions before the beginning of the November 2013 outburst. We have combined these observations in order to determine if SXP 5.05 has been previously detected. These data give a combined exposure of $\sim 19 \text{ ks}$ of *Proportional Counter* (PC) mode data, taken between 2006 November and 2010 September. We find no point source inside the *INTEGRAL* error circle, and calculate a 3σ upper limit of $7 \times 10^{-14} \text{ erg s}^{-1} \text{ cm}^{-2}$ (0.5 – 10 keV), constraining the pre-outburst flux to less than $3 \times 10^{34} \text{ erg s}^{-1}$, assuming a standard SMC distance of 60 kpc, or less than 0.02% Eddington luminosity assuming a $1.4 M_{\odot}$ neutron star compact object.

2.3.2 *Swift* detection and localization of SXP 5.05

The first *Swift* post-outburst observation of SXP 5.05 started on MJD 56601 (2013 November 5) at 23:42UT, with a 1 ks PC mode observation. Inside the *INTEGRAL* error circle, we found a single bright point source, and assumed to be the X-ray counterpart of SXP 5.05.

We determined the position of SXP 5.05 utilizing the methods described by Goad et al. (2007) and Evans et al. (2009). We find that the coordinates of SXP 5.05 are RA(J2000) = $00^{\text{h}}57^{\text{m}}02^{\text{s}}.34$, Dec(J2000) = $-72^{\circ}25'55''.34$, with an estimated uncertainty of $2.6''$ radius (90% confidence).

2.3.3 *Swift* Timing Analysis

For analysis of the pulsar timing of SXP 5.05, we utilize only data taken using the XRT ; Windowed Timing (WT) mode, which provides high time resolution (2.2ms) event time-tagging. Photon Counting (PC) mode data in this case

are not useful as they suffer both from pile-up and aliasing issues, as the PC mode frame time (2.506s), is very close to 50% of the mean pulsar period of SXP 5.05. Exposure times are typically between 3–6 ks each.

All *Swift* XRT timing analysis was performed using the HEASoft 6.14. The individual downloaded datasets for each observation were first reprocessed through *xrtpipeline* with the coordinates for SXP 5.05 given, and then event files were extracted using *xselect*, utilizing a circular extraction region of 20 pixels radius centered on the coordinates of SXP 5.05. A barycentric correction was then applied using *barycorr*.

Period searching was performed using a Z_2^2 search (Buccheri et al, 1983), with the period search centered on $P = 5.05 \text{ s}$, a step size of 10^{-5} s , over 10000 steps, searching a period range of 5.0 – 5.1 s. The reported period for each observation corresponds to the peak of the Z_2^2 distribution. In order to double check the determined period we also ran the period search with *efsearch*, which utilizes a χ^2 folding search. We found the results to be consistent, although the Z_2^2 search proved to be more sensitive to detecting the period where the total counts in the exposure were low. The detected periodicities are plotted in Figure 9.

In order to estimate errors on the measured period, we implemented a Monte Carlo approach similar to that described by Gotthelf, Vasisht & Dotani (1999). Based on the folded light-curve of the WT event data, we create a simulated event file with the same total event rate and average event rate as the *Swift* observations. These are then period searched in the same way as above. This process is repeated 1000 times, and the resultant scatter in the measurement of the period in these simulated datasets is used to derive the error on the measured period. We found that in all cases the errors were larger than the 10^{-5} s step size used in the period search, showing that no benefit would be obtained from searching with a higher time step.

We note that after 2014 March 4, the periodicity could no longer be detected with sufficient accuracy due to the dropping brightness of SXP 5.05.

The orbital motion of a neutron star about its optical companion presents a window through which to study the orbital parameters of that binary system. In Be/X-ray binaries, the orbital motion is most typically seen through the Doppler shifting of the X-ray light emitted by the pole of the neutron star, as the neutron star travels along its orbit. By decoupling the change in this period caused by the orbital motion from the changes due to the accretion torques exerted on the neutron star, we can model this motion and determine the system parameters. This method has been used extensively in the Milky Way to calculate these parameters for several HMXBs, and more recently in the SMC to find the parameters for the first group of systems outside the Galaxy (Townsend et al., 2011a, 2011b).

The details of the *Swift*/XRT measured periods are presented in Table 2.3.3.

2.3.4 *Binary solutions*

As the detected spin period will be a function of the orbital parameters, accretion rate and inclination of the system, one needs to be careful in extracting out the required information. In the simplest example of the orbital plane being

Table 5. Journal of *Swift*/XRT observations done in Windows Timing mode and the determined pulse period for each observation.

Observation Date (mid-point MJD)	Observation Span (days)	Period (s)	Exposure (s)	ObsID
56640.01981	0.10045	5.05006 ± 0.00004	7161	00033038021
56671.34285	0.34258	5.04791 ± 0.00001	3822	00033038089
56673.20578	0.0736	5.04868 ± 0.00008	4032	00033038090
56675.53490	0.07025	5.05037 ± 0.00006	3707	00033038091
56677.27352	0.0711	5.05219 ± 0.00005	3780	00033038092
56679.47241	0.07288	5.05387 ± 0.00004	3758	00033038093
56681.27895	0.14383	5.05207 ± 0.00155	3274	00033038094
56683.77892	0.16333	5.04992 ± 0.00003	3552	00033038095
56685.07323	0.07008	5.04873 ± 0.00007	3408	00033038096
56687.14618	0.07538	5.04776 ± 0.00004	3395	00033038097
56689.20835	0.20087	5.04789 ± 0.00003	3900	00033038098
56693.21210	0.20254	5.05070 ± 0.00002	4918	00033038100
56694.50240	0.30328	5.05192 ± 0.00002	6032	00033038102
56696.07214	0.34128	5.05342 ± 0.00002	4330	00033038103
56697.68006	0.46938	5.05291 ± 0.00001	4738	00033038104
56699.18124	0.10148	5.05196 ± 0.00009	2922	00033038105
56701.34264	0.20489	5.04927 ± 0.00004	4264	00033038106

perpendicular to our line-of-sight (inclination, $i = 0$), any change seen in the spin period of the neutron star should be due to torque produced by accretion onto the neutron star surface. The most common result of this is for the spin period to decrease over the length of the outburst. In the more generic case of the orbital plane being at some angle to our line-of-sight, we see more complicated spin period variations. Convolved with the standard spin-up of the neutron star, there will be a shifting of the X-ray pulse arrival times caused by the orbital motion of the neutron star around the counterpart. It is for this reason that our orbit fitting routine simultaneously fits a simple spin-up model and an orbital radial velocity model.

$$P_{obs} = \left(1 + \frac{v(t)}{c}\right) P(t) \quad (2)$$

where P_{obs} is the detected spin period of the pulsar, c is the speed of light and $v(t)$ is the binary radial velocity. $v(t)$ is calculated using the IDL routine `BINRADVEL.PRO` from the `AITLIB`² which uses the procedure of Hilditch et al. (2001) and employs the method of Mikola (1987) to solve Kepler's equation. The spin of the pulsar is given by $P(t)$ which is defined as

$$P(t) = P(t_0) + \dot{P}(t - t_0) - \frac{\ddot{P}(t - t_0)^2}{2} \quad (3)$$

where \dot{P} and \ddot{P} are the spin-up and change in spin-up of the neutron star. We perform a least-squares fit on the data and return the best fitting parameters. The full spin-up component shown in Equation 3 was fit to the data shown in Figure 9. One can see from this figure that there is a period between MJD 56670–56700 during which the source was densely monitored. On top of this main data set, there are a few scattered data points over the preceding 30 days. In trying to fit the

Table 6. The orbital parameters for SXP 5.05 from the bootstrapping analysis of *Swift* XRT data.

Parameter		Orbital Solution
Orbital period	$P_{orbital}$ (d)	17.13 ± 0.14
Projected semimajor axis	$a_x \sin i$ (light-s)	142.4 ± 2.5
Longitude of periastron	ω ($^\circ$)	19.7 ± 4.5
Eccentricity	e	0.155 ± 0.018
Orbital epoch	$\tau_{periastron}$ (MJD)	56680.45 ± 0.22
First derivative of P	\dot{P} (10^{-10}ss^{-1})	-2.25 ± 0.30

above model to the full data set, we found that we could not obtain a satisfactory fit. A possible explanation for this is that the accretion rate has changed markedly over the 2 months, as shown in Figure 1, meaning there is some higher order period change that we cannot model. The \dot{P} component can be estimated from the luminosity changes in Figure 1 and turns out to be roughly $10^{-17} \text{s}^2/\text{s}$. However, when fitting this component in the model, the parameter always tended to zero. Forcing the parameter to be at this value resulted in even worse fits. To account for this, we decided to only model the data from MJD 56670 (2014 Jan 13) onwards as the effect from the change in flux would be minimised and we would still keep most of the data. We thus removed the \ddot{P} component from the model and only fit a \dot{P} to minimise the number of variables. The result of this was a much more stable fit, with a better χ^2 value. However, the overall χ^2_ν was still unacceptable, forcing a more thorough investigation into the binary model.

The reason for a poor fit using the above method seems to result from some random change in the spin period of the neutron star on top of the expected orbital and spin-up signatures. We believe this may arise from a highly variable accretion rate caused by a clumpy wind. Such over-densities in an otherwise uniform wind, would result in unpredictable changes in the accretion rate and, hence, the spin period measured. As this is almost impossible to incorporate into our model, we cannot be sure that the measured period and associated errors can be well described by the model in Equation 2. To account for this, we decided to investigate which data points might be affected by performing a bootstrapping with replacement simulation. The method involved remov-

² Developed at the Institut für Astronomie und Astrophysik, Abteilung Astronomie, of the University of Tübingen and at the University of Warwick, UK; see <http://astro.uni-tuebingen.de/software/idl/>

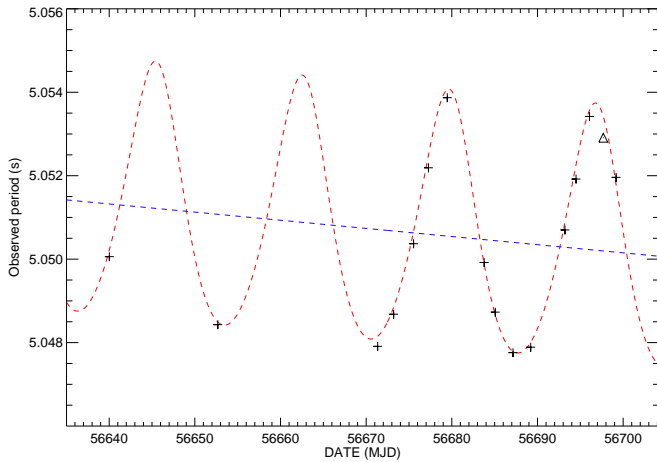


Figure 9. The spin period of SXP 5.05 as measured by *Swift*. Both the orbital motion of the neutron star and the accretion driven spin-up are visible. Overplotted are the spin-up and orbital model from a single fit to the full dataset. It is therefore not the final solution presented in Table 5, but is to be used as an indication of the data used in the fit and the approximate shape and magnitude of the model components. The one point shown as a triangle was not included in the final fit - see text. We note the independent pulse period determination from *XMM-Newton* of 5.0484s on MJD56652 is in complete agreement with the *Swift* measurement of the same date.

ing a variable number of random points from our dataset and fitting the above model each time to the subset generated, with the expected result being a normal distribution of results for each of the binary parameters. The total number of points in our dataset was 16 and the number of variables 7. The small number of degrees of freedom meant that our distributions were not perfectly Gaussian, but had small bumps and shoulders in some of the parameter distributions. An example distribution and Gaussian fit can be seen in Figure 10. The number of iterations was also kept quite high (100000 runs) to try and mediate this problem. During the initial run, we found one of the 16 data points was causing a poor fit when it was included much more often than the inclusion of the other points. We thus excluded this point (the penultimate one) from the fit and ran the bootstrap again. These results were much more satisfactory and we were able to fit a single Gaussian to all the distributions, to obtain final values and errors. These are presented in Table 6. The parameter errors are likely quite conservative based on the bootstrapping method used as opposed to a single reduced Chi-squared fit, but rightly so given the difficulty had in fitting these data. The system parameters are discussed further in the latter sections of this paper.

2.3.5 *Swift* spectral pulse profile information

Shown in Figure 11 is the pulse fraction defined as in Equation 1 as a function of the general source brightness. This figure reveals that the typical pulse fraction in the *Swift* 0.5–10 keV energy band is in the 20–40% range, with strong evidence for a correlation between the pulse fraction and the luminosity.

Figure 12 reveals more details of these pulse profile changes with luminosity. In this figure the decreasing com-

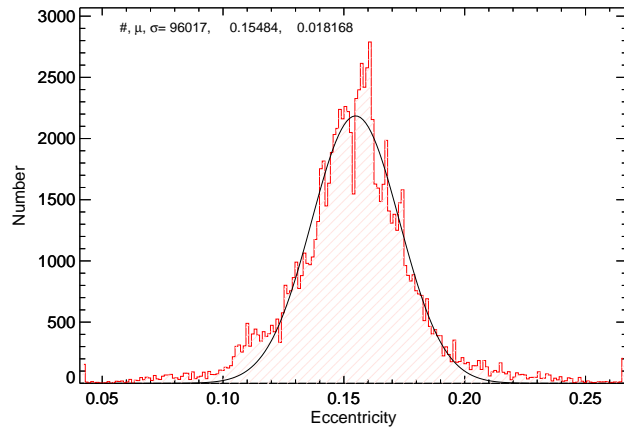


Figure 10. Example distribution from our bootstrapping fitting method. We show the eccentricity parameter, though it is representative of the other parameters. The histogram shows the distribution of the 96,000 fits in the simulation and the fitted Gaussian is overlaid with the mean and error stated.

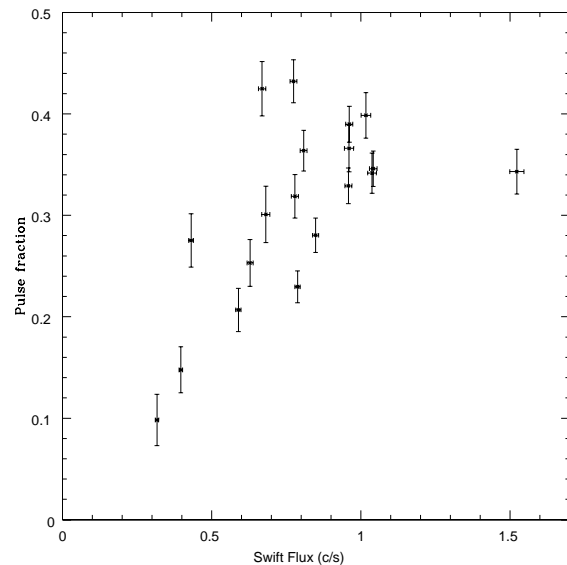


Figure 11. *Swift*/XRT pulse fraction versus 0.5–10 keV X-ray count rate.

plexity of the pulse profile is revealed as the source moved from a bright state (~ 1.5 counts/s) to a fainter state (~ 0.5 counts/s). Similar structural changes, but with better signal-to-noise ratios, were reported for EXO 2030+375 by Parmar, White & Stella (1989). Those authors attributed the changing pulse profiles to be related to a switch in the dominant emission profile from a fan-beam to a pencil-beam as the luminosity changed. Specifically they observed a lack of sharp features in the pulse profile at high luminosity in EXO 20230+375, with the profiles becoming more complex as the luminosity fell from $\sim 1 \times 10^{38}$ erg/s to $\sim 1 \times 10^{37}$ erg/s. Below this flux level the profiles smooth out again, though ultimately the diminishing signal-to-noise ratio as the flux

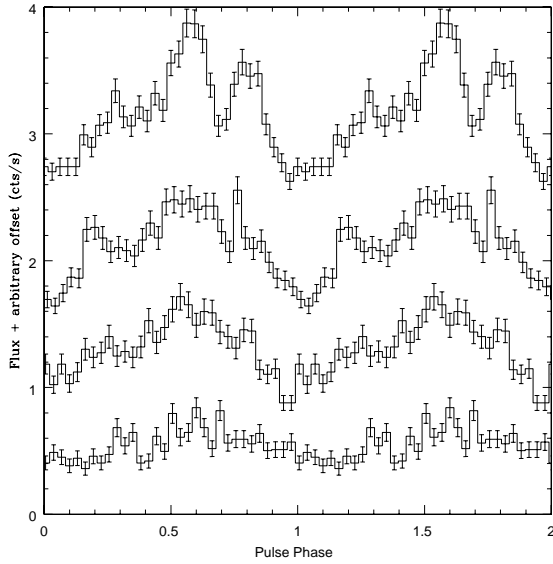


Figure 12. *Swift*/XRT pulse profiles (0.5–10 keV). From the top, the profiles are arranged in decreasing source brightness and show observations made on MJD 56639.9, MJD 56671.0, MJD 56683.6 & MJD 566701.1; the profiles have been separated by adding 2.0, 1.2, 0.5 & 0 cts/s respectively to their count rates. In addition, the profiles have simply been lined up in phase by eye.

approaches $\sim 1 \times 10^{36}$ erg/s makes this somewhat uncertain.

The brightest *Swift* profile shown in Figure 12 corresponds to a source luminosity of $\sim 2 \times 10^{37}$ erg/s and reveals a strong pulse and inter-pulse peak separated by ~ 0.2 in phase. This flux level corresponds to that of maximum pulse shape complexity in EXO 2030+375 and hence may support the idea of a combination of pencil and fan beam emissions at this stage. Similar profiles were seen in the *XMM* observations - see Figure 6. More detailed modelling, beyond the scope of this paper, including the possibility of an off-set dipole field as suggested by Parmar, White & Stella (1989) for EXO 2030+375, is required to understand the pulse profile shapes seen here in SXP 5.05.

2.4 RXTE

The SMC was observed on an almost-weekly basis using the *Rossi X-ray Timing Explorer* (RXTE). The results of searching through these ~ 13 years worth of observations of the SMC reveals that SXP 5.05 was frequently within the field of view, often at a high collimator response, but was only detected 5 times at a confidence rating of $\geq 99\%$ (see Galache et al, 2008 for an explanation of the RXTE data analysis techniques). The dates detected and the associated periods are presented in Table 7. Using the orbital parameters presented in Table 6 reveals no obvious phase pattern for the detections, though the uncertainty of 0.14d on the orbital period means that extrapolating back over a decade makes the phase value increasingly uncertain. Typical errors on a 5-10s pulsar detected by RXTE at the 99% confidence level are ≤ 0.01 s (Galache et al, 2008). The scatter in pulse periods around 5.02-5.09s suggests that the RXTE data are revealing

Table 7. Historical RXTE detections of SXP 5.05 at $\geq 99\%$ confidence.

Date (MJD)	Collimator Response	Period detected (s)
55179.3	0.89	5.02
52458.6	0.26	5.06
52882.6	0.78	5.07
53487.9	0.78	5.09
55785.7	0.89	5.06

Table 8. OGLE identifications for the optical counterpart to SXP 5.05

OGLE project phase	I-band ID	V-band ID
OGLE II	smc sc7.266703	-
OGLE III	smc 108.8.26225	smc108.8.v.30808
OGLE IV	smc 719.18.378	smc719.18.v.340

a pulsar close to equilibrium spin period outside of the times of major outbursts. This would be consistent with the overall behaviour of SMC Be/X-ray binaries seen in the SMC (Klus et al, 2014).

3 OPTICAL OBSERVATIONS

3.1 OGLE

Optical Gravitational Lensing Experiment data (OGLE-II, III and IV, Udalski et al. 1997; Udalski 2003) were used to investigate the long-term behaviour of SXP 5.05. OGLE has been regularly monitoring this object since 1997 with the 1.3-m Warsaw telescope at Las Campanas Observatory, Chile, equipped with three generations of CCD camera: a single $2k \times 2k$ chip operated in driftscan mode (OGLE-II), an eight chip 64 Mpixel mosaic (OGLE-III) and finally a 32-chip 256 Mpixel mosaic (OGLE-IV). Observations were collected in the standard I-band.

OGLE photometric I-band observations of SXP 5.05 were obtained from MJD 50627(1997 June) up to MJD 56689 (2014 January). The OGLE source identification in each phase is given in Table 8.

The full photometry from this 16-17 year coverage is shown in Figure 13. From this figure it is immediately apparent that the source exhibits considerable rapid optical variability on timescales of months to years, culminating in the recent rise of 0.3 magnitudes within a month. Not obvious in this figure is the ~ 17 d modulation seen by close examination of the decline from outburst which occurred at the start and end of this whole data set. This was initially commented upon by Schmidtke & Cowley (2013) and is shown very clearly in Figure 14 during the current outburst.

Shown in Figure 15 are the trailed Lomb-Scargle power spectra obtained from each annual sample of OGLE data. The date were detrended and processed using the techniques described in Bird et al. (2012). It is immediately apparent that not only is there power at the fundamental frequency at the start and end of the data run, but in between the main power peaks are often at one of the harmonics. Adding together all the different peaks from all the years gives a final value of 17.13 ± 0.01 d for the fundamental optical period.

Shown in Figure 16 are the observed (V-I) colours versus the I-band brightness. It is clear there is an obvious cor-

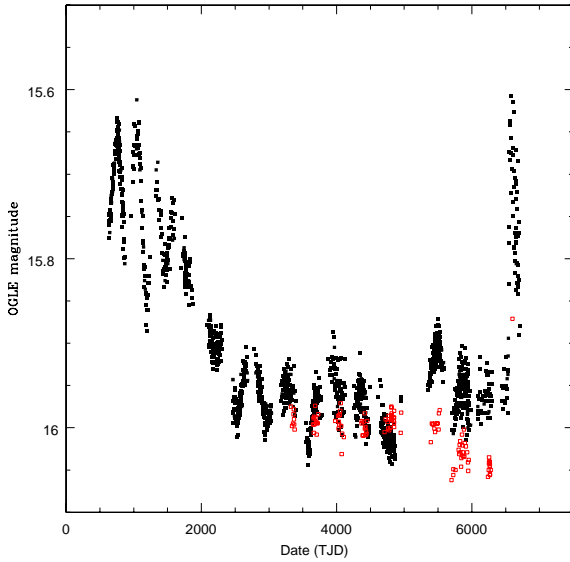


Figure 13. OGLE monitoring of SXP 5.05 during the period 1997-2014. The solid black symbols are I-band data, whereas the open red symbols are V-band data. TJD = JD - 2440000.5.

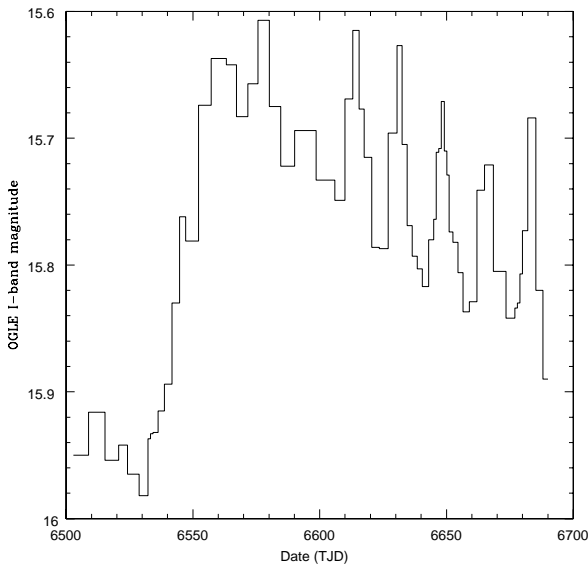


Figure 14. OGLE monitoring of SXP 5.05 during the 2013-2014 outburst. The ~ 17 d modulation is very clear. TJD = JD - 2440000.5.

relation between these two parameters as seen in other similar systems (e.g. Vasilopoulos et al. 2014, Coe et al. 2012). The V band data were taken whilst the source was not in X-ray outburst and so this figure shows the general fluctuations in the circumstellar disk that occur most of the time. To check whether the colour variations have anything to do with the neutron star orbit, the colour data are plotted versus the 17.13d X-ray ephemeris in the lower part of Figure 16. There appears to be no relationship with this optical phase and therefore one can conclude that in the quiescent state the neutron star is not exerting any influence over any disk pa-

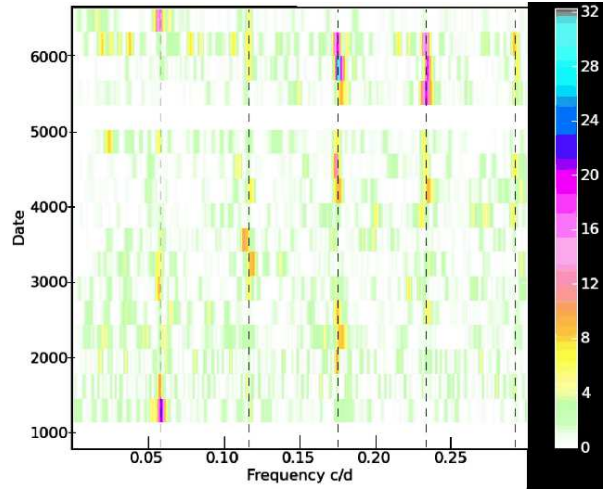


Figure 15. Period searching through OGLE data - the fundamental is at a frequency of 0.058 c/d and is indicated by the left-most dashed vertical line. The other vertical dashed lines show the position of various harmonics of the fundamental. The colour scale represents the Lomb-Scargle peak power value.

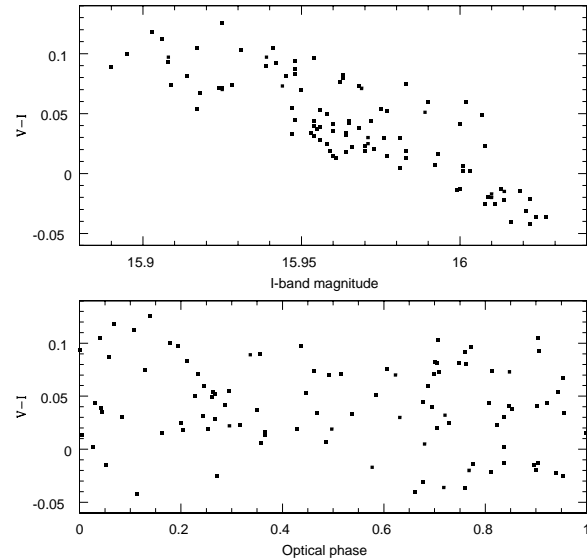


Figure 16. OGLE observed colours versus brightness (top panel), and colours versus optical phase (lower panel).

rameters that could result in changing temperatures. However, given the uncertainties on the orbital period quoted in Table 1 there is considerable uncertainty in the phases of the early OGLE data - but even selecting just the most recent few years does not change the lack of pattern in Figure 16.

This is in contrast to that seen in LXP 8.04 where a strong colour effect was seen correlated with the optical phase (Vasilopoulos et al. 2014). Unfortunately there are not enough V band observations during the outbursts of SXP 5.05 just to select those epochs for a more detailed study of the larger disk behaviour.

3.2 SALT

SXP 5.05 was observed with the Robert Stobie Spectrograph (RSS) (Burgh et al. 2003, Kobulnicky et al. 2003) on the Southern African Large Telescope (SALT) at various epochs during outburst with a view to spectral classification of the donor star and $H\alpha$ monitoring of the Be star circumstellar disk. An observation log is presented in Table 9.

The spectra were corrected for bias and cross talk by the SALT data reduction pipeline (Crawford et al. 2010). Subsequent data reduction was performed in IRAF³. The two-dimensional images were wavelength calibrated using observations of Neon (PG1800) or Copper Argon arc lamps (PG2300) and then rectified before one dimensional, background subtracted spectra were extracted. Grating PG2300 spectra were flatfielded.

Figure 17 shows the continuum-normalised, smoothed SALT spectrum observed on 2013-11-02 with grating PG2300, blue shifted by the recession velocity of the SMC (150 km s^{-1}). The Balmer lines and He I lines characteristic of early type stars are clearly visible. He II at 4541 and 4200 are detected, while $H\beta$ is in emission, which places the spectral type at B0e–B0.2e, according to the classification criteria presented in Evans et al. (2004). SXP 5.05, which appears to be a member of NGC 330, has been previously observed as part of the VLT-FLAMES survey of massive stars (Evans et al. 2006). The spectral type determined by those authors is B0.5e. It is hard to reconcile the SALT data with this slightly later type, but we do note that our result implies a significant discrepancy between the mass determined from the spectral type ($\sim 16.0 M_{\odot}$) and the dynamical mass estimates ($\sim 13.0 M_{\odot}$) discussed in Section 4.1 below. This system is the first one in the SMC to have a mass determination from both spectral type and dynamical methods and it reveals a substantial gap between the two methods.

The spectrum at that time (between 2003 July and 2004 January) shows very clearly the He II 4686 in absorption consistent with spectral types as cool as B0.5. Subsequent spectra observed with SALT on 2013-11-11 and 2013-12-05 confirm the He II 4686 line in emission, and we interpret this as a feature arising from an accretion disk around the neutron star.

Shown in Figure 18 are the $H\alpha$ profiles. The profiles start off with a clear triple peak structure, then evolves to a double peak structure by MJD 56631 (2013 December 5), but the triple peak structure seems to return by the last observation MJD 56649 (2013 Dec 23). 3 Gaussian lines were fitted to all observations. The central component is very narrow (sigma $\sim 1 \text{ \AA}$) and located at the $H\alpha$ rest wavelength for the SMC, while the two outer components are broader. It is worth noting that the VLT-Flames spectra from Evans et al. (2006) also show this triple structure in the $H\alpha$ emission line.

The equivalent width is statistically consistent with a mean value of $-6.9 \pm 0.2 \text{ \AA}$ throughout the observations. Because the SALT coverage is relatively sparse it is not possible to say with any certainty if there is any correlation between the $H\alpha$ EW values and the phase of the OGLE I-band modulation. However, it is perhaps worth noting that the 4 largest

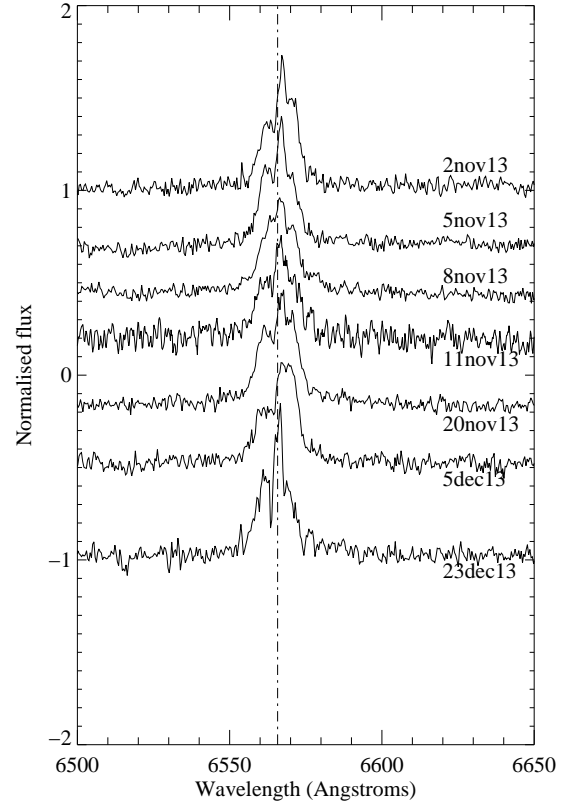


Figure 18. $H\alpha$ 6562.3 \AA emission line profiles. Spectra have been rectified with arbitrary offsets on the y-axis to show the line profile evolution with time.

$H\alpha$ EW values correspond to peaks in the I-band modulation, and the 3 weakest to the baseline flux level. This suggests that the interaction of the neutron star with the circumstellar disk at periastron creates some disturbance of the disk structure.

The double-peaked $H\alpha$ profile can be used to calculate the radius of the circumstellar disk (Coe et al., 2006). Using the average peak separation of $8 \pm 2 \text{ \AA}$, and making the assumptions that the circumstellar disk lies in our line of sight ($i = 90^\circ$), and the true mass of the star is $13.0 M_{\odot}$. The radius is given by:

$$R_{\text{circ}} = \frac{GM \sin^2 i}{(0.5 \Delta V)^2} \quad (4)$$

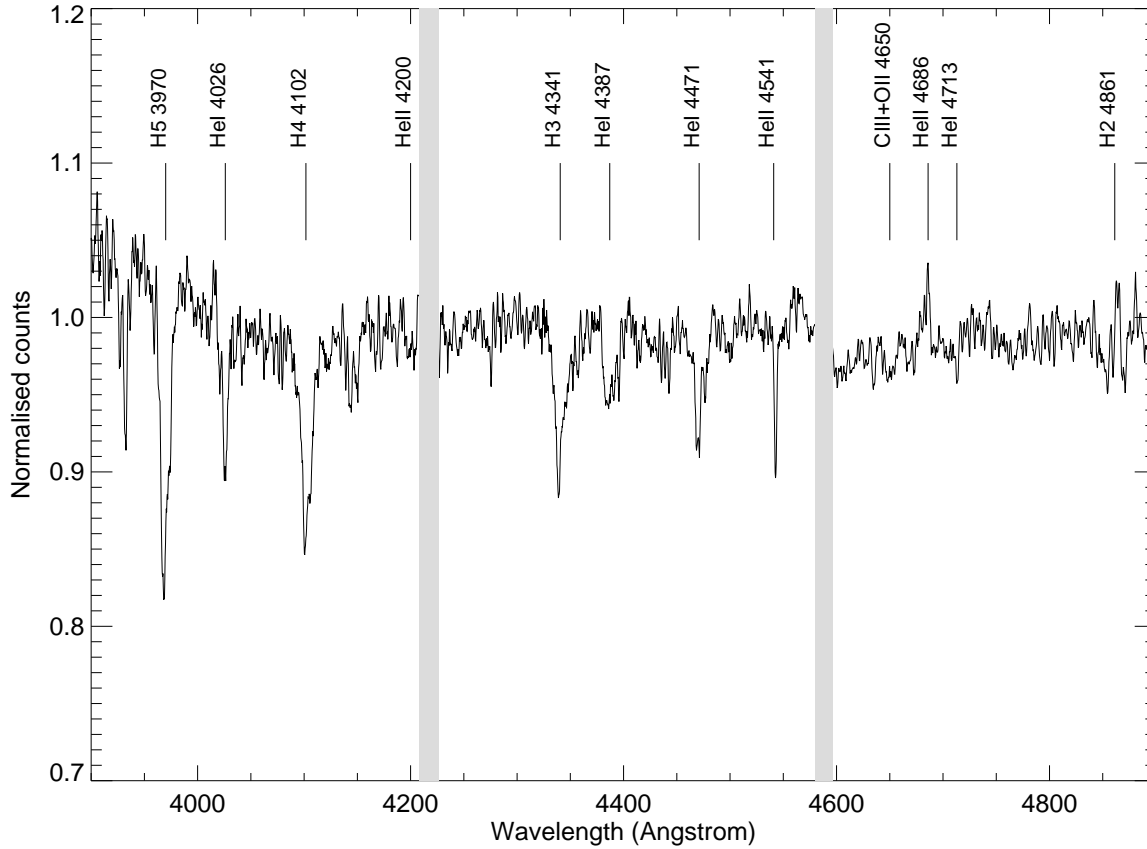
where ΔV is the peak separation in velocity space. This results in a radius of $(2.6 \pm 1.8) \times 10^{10} \text{ m}$ (clearly it can be smaller if the inclination is different).

Another approach to estimating the size of the circumstellar disk is simply to use the $H\alpha$ EW. Assuming the mean value of $-6.9 \pm 0.2 \text{ \AA}$, the size of the B0V star, and utilising the relationship in Zamanov et al. (2001), the radius of the disk is found to be $(4.8 \pm 0.1) \times 10^{10} \text{ m}$ - just about in agreement with the previous number.

³ IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy (AURA) under cooperative agreement with the National Science Foundation

Table 9. Journal of *SALT* spectroscopic observations of SXP 5.05

Date (MJD)	Grating	Exposure (s)	Resolution (Å)	Coverage	H α EW (Å)
56598	PG2300	1000	2.0	3800–4900 ÅÅ	–
56598	PG1800	840	4.7	5950–7250 ÅÅ	-6.6 (0.5)
56601	PG1800	840	4.7	5950–7250 ÅÅ	-7.3 (0.5)
56604	PG1800	840	4.7	5950–7250 ÅÅ	-6.6 (0.5)
56607	PG2300	2×1220	2.0	3800–4900 ÅÅ	–
56607	PG1800	840	4.7	5950–7250 ÅÅ	-5.6 (0.5)
56616	PG1800	840	4.7	5950–7250 ÅÅ	-8.0 (0.5)
56631	PG2300	1000	2.0	3800–4900 ÅÅ	–
56631	PG1800	840	4.7	5950–7250 ÅÅ	-6.8 (0.5)
56649	PG1800	1000	4.7	5950–7250 ÅÅ	-7.3 (0.5)

**Figure 17.** The optical spectrum of SXP 5.05, with lines of interest overplotted. The grey shaded areas are the chip gaps in the detector.

4 DISCUSSION AND CONCLUSIONS

4.1 Binary models

The binary mass function is described by the following

$$f_X(M) = \frac{4\pi^2 a_x^3 \sin^3 i}{GP^2} = \frac{M_c^3 \sin^3 i}{(M_x + M_c)^2} \quad (5)$$

where M_x and M_c are the masses of the neutron star and the optical counterpart, i is the inclination of the orbital plane to the line of sight, P is the orbital period, a_x is the semi-major axis of the ellipse travelled by the neutron star and e is the eccentricity. Thus, if the mass functions of both stars are known, along with the inclination angle, one can dynam-

ically calculate the masses of the two stars. Substituting in the values from Section 2.3.4, we obtain a mass function of $10.56 M_\odot$. This is quite high relative to many other Be/X-ray binary systems, but not as large as most supergiant systems (see Townsend et al. (2011) for a summary of known binary parameters and mass functions).

It can be seen in Figure 19 that the time during the eclipse when the flux drops close to zero is at MJD 56648.6. Given the 17.13 d orbital period found in the OGLE data this point is directly behind the companion star as viewed from Earth. This observation supports our theory that the eclipse (at least the deepest part of the occultation) is caused by the companion. If this is correct, the inclination of this system is almost edge on to our line-of-sight. Taking an orbital phase of

0.04 as the point of eclipse, the separation of the neutron star and the Be star may be determined to be 4.08×10^{10} m. Combining this with a stellar radius of B0.2V star of 5.2×10^9 m gives a minimum orbital inclination of $i = 82.7^\circ$.

Constraining the inclination in this way can give us an idea of the actual orbital size, the orbit disk mis-alignment, the actual geometry of the binary system and, possibly, the mass of the two bodies. This latter point is very interesting as it would be the first time a neutron star mass measurement will have been dynamically determined in a Be/X-ray binary - however it will require future radial velocity observations of the optical star to determine this unambiguously. Even without this radial velocity curve, we can place constraints on the masses of the two bodies, assuming $i = 90^\circ$. Given a set of realistic to extreme masses for the neutron star, say, 0.5, 1.0, 1.4, 2.0 and $2.5 M_\odot$, we can solve the binary mass function for M_c . The resulting solutions are $M_c = 11.5, 12.3, 13.0, 13.8$ and $14.5 M_\odot$ using the respective neutron star masses above. Reducing the inclination to $i = 80^\circ$ only changes these values by $\leq 5\%$. Thus, even with an extremely large neutron star mass of $2\text{--}2.5 M_\odot$, the mass of the counterpart is $2\text{--}3 M_\odot$ less than standard mass of a B0.2V star which is around $16 M_\odot$. Discrepancies of this kind between stellar classification and observed luminosities plus inferred masses & temperatures, have previously been reported for HMXB companions by Conti (1978) and Kaper (2001). This can only be explored accurately with eclipsing systems, so SXP 5.05 will be a valuable addition to the small number known of such systems.

Assuming an inclination of $i = 90^\circ$, the actual size of the orbit is given in Table 6, and is 4.27×10^7 km. In Section 3.2 the size of the circumstellar disk was estimated to be $2.6\text{--}4.8 \times 10^7$ km, which is similar to the orbital radius and confirms that the neutron star is in contact with the disk during this outburst. This is discussed further below.

The optical flare seen regularly in SXP 5.05 is at a phase of 0.15 relative to periastron passage according to the orbital model. Though delays between optical flares and periastron passages are well established in Be/X-ray systems, this is much later than any other examples we have in the SMC. If we use the X-ray ephemerides from the sources observed as part of the 12 years of observations of the SMC using RXTE (Galache et al, 2008) and compare these to the optical ephemerides determined by Bird et al. (2012) we can derive both optical and X-ray epochs of maximum amplitude and compared them. These results may then be converted into a phase delay for that system based on the optical orbital period - see Table 10. Clearly SXP 5.05 exhibits a much greater phase delay between X-ray & optical outburst than any of these systems. If we translate the phase delays given in Table 10 into elapsed time, then values in the range 2-13d are determined. These values of a few days may well be indicative of some physical characteristic of the circumstellar disk response times within the accretion process.

In addition, we note that the phase of the optical outbursts coincides with that of the X-ray eclipse. We have no explanation for this alignment except to invoke coincidence.

4.2 Modelling the circumstellar disk

Shown in Figure 19 is a detailed study of the *Swift*/*XRT* flux during the occultation event for which there are the most ob-

Table 10. Phase delays between optical outbursts and the X-ray outburst for a small sample of SXP systems.

Source	Phase delay	Orbital period(d)
SXP 7.78	0.04	44.8
SXP 46.6	0.03–0.10	137.4
SXP 91.1	0.04	88.0
SXP 756	0.00–0.01	394.0

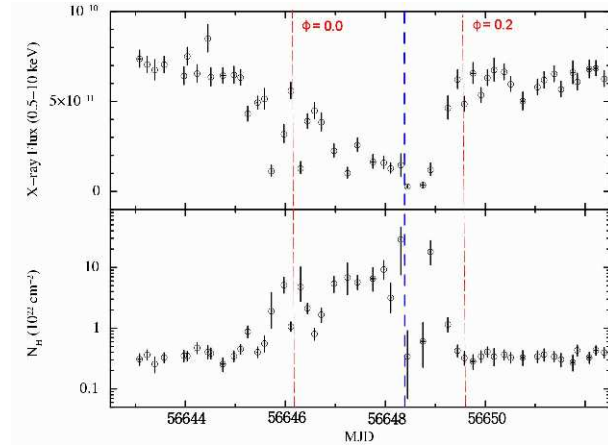


Figure 19. Variations in the *Swift*/*XRT* measured X-ray flux and absorption during the occultation/eclipse. Note that the upper panel has a linear vertical scale, whereas the lower panel is logarithmic. The vertical red dashed lines show the binary phase, and the vertical dashed blue line indicates the binary phase (not the date) of the first *XMM-Newton* observation..

servations. From this plot it can be seen that the suppression of the source flux occurred gradually over ~ 120 hours, but in a somewhat erratic manner. After the point of minimum flux, there was a more rapid recovery phase. The total size of the occultation event greatly exceeds the simple effect of eclipsing the neutron star with its companion Be star.

The asymmetry in the X-ray absorption is hard to explain with a circular circumstellar disk in the plane of the neutron star orbit, instead we invoke the possibility of the disk being in a plane orthogonal to the plane of the neutron star orbit. Then, depending upon the orientation of the disk, it is possible to create asymmetries in the absorbing column to the neutron star.

We have attempted to model the absorption profile between phases 0.7 and 0.25 by considering the attenuation of the neutron star emission by the Be star's circumstellar disk. Attenuation by the disk is essential to produce the broad orbital effect seen, which is incompatible with an eclipse by the Be star alone. We require a disk that is partially orthogonal to the orbital plane of the binary, such that the neutron star can move 'behind' the disk with respect to our line of sight for a significant phase range. In order to induce the observed asymmetry in the absorption, we further require that the line of sight is close to the plane of the circumstellar disk, so that the neutron star is observed obliquely through the disk at all times. We have implemented a 2D disk model that is able to capture the key geometric features needed - a truncated, flared disk with a radial density fall-off. We assume the power-law radial profile and simplified disk scale height model from Rivinius, Carciofi & Martayan (2013), with a linear growth in scale height reaching $3.5R^*$ at $20R^*$. We imple-

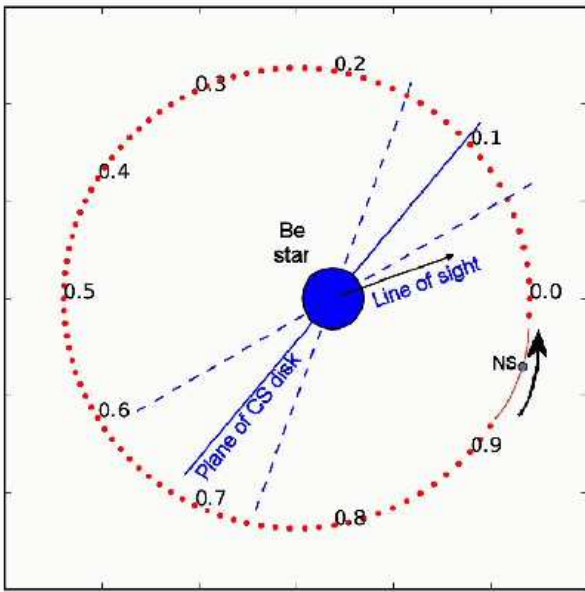


Figure 20. The orientation of SXP 5.05 system with respect to our line of sight. Phase 0.0 represents periastron. The plane of the circumstellar disk is orthogonal to the neutron star orbital plane. The blue dashed lines represent the outline of the flared circumstellar disk used in modelling the X-ray data. The solid black arrow represents our line of sight with respect to periastron ($\theta_{los} = 20^\circ$). The dotted black line ($\theta_{disk} = 50^\circ$) represents the central axis of the circumstellar disk which is inclined orthogonally to the plane of the orbit.

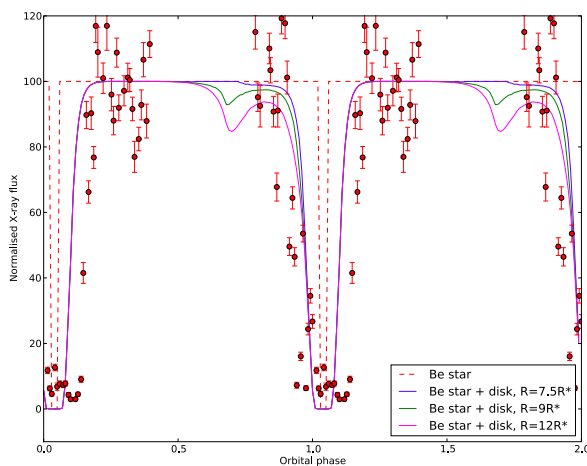


Figure 21. Implications of the current model for different disk sizes compared to the *Swift*/XRT data. The red dashed line shows the effect of a simple geometric eclipse by the Be star.

ment a truncated circumstellar disk by imposing a hard density drop to zero at the desired distance from the Be star.

With reference to Figure 20 we can see the neutron star will be obscured by the disk between phases of 0.7 and 0.2, assuming the black arrow represents our line of sight ($\theta_{los} = 50^\circ$) and the dotted line ($\theta_{disk} = 70^\circ$) represents the

central axis of the circumstellar disk. These values are only approximate, as the disk scale height prevents any meaningful discussion of the ‘thickness’ of the disk, and the flaring shown is only for illustration. Any eclipse by the Be star, if present, would be at phase 0.04. For each phase of the neutron star orbit, the cumulative disk density along the line of sight between the neutron star and observer is calculated, and then converted into an attenuation value. Initially (at phase 0.7) the obscuration is low as the neutron star is only obscured by the low-density outer regions of the flared disk. But then the obscuration increases slowly as the line of sight is through the full thickness of the outer disk. The attenuation then rises rapidly towards phase 0 as the line of sight moves towards the dense central regions of the disk. A much sharper recovery is seen between phase 0.1 and 0.2, as the neutron star traverses the disk again.

From Figure 1 it can be seen that this occultation event occurred against a general background of the X-ray flux declining globally in a very linear manner between MJD 56635 and the end of the observations. Apart from the occultation event the flux can be accurately described by the simple empirical relationship over this time interval - see Equation 6.

$$F(t) = -0.01667t + 945.59 \quad (6)$$

where $F(t)$ is the *Swift* /XRT count rate, and t is the date in MJD.

Figure 21 shows a family of curves for a power-law radial profile with an index of $n=3.75$ and a base density of $\rho_{zero} = 10^{-10} \text{ g/cc}$ (within the suggested range of Rivinius, Carciofi & Martayan (2013), but definitely towards the high end, implying that this is a dense disk). The extremes of the curves have been normalised to fit the *Swift* /XRT count rates scaled according to Equation 6. Three different disk truncation radii are simulated, and the effect on the early stages of obscuration (phase 0.7) are obvious. More extended disks also produce higher attenuation around phase 0.0, as the line of sight is looking through the disk, although the low density of the outer disk makes this a small effect.

Overall, we conclude that even a simple geometric model of this type is able to explain some of the key features seen in the folded x-ray light curve, principally the long duration over which obscuration is seen, and the asymmetry of the obscuration around the predicted eclipse point. However, we emphasise that this is in no way a parameterised fit to the observed data. Furthermore it is clear that some parts of the model fit do not describe the data very well - specifically the egress phase. Many of the parameters are quite degenerate (notably the disk density and the radial density profile) and we can only state that the ‘typical’ disk parameters given by Rivinius, Carciofi & Martayan (2013) do result in a reasonable attenuation profile.

Perhaps the biggest difficulty faced by this simple model is explaining in detail the obscuration observed between phases 0.9 and 1.0. Here we observe flux variations by a factor of 2 in very short timescales which obviously cannot be reproduced by the smooth disk model used here. The presence of such rapid changes in attenuation implies the disk is not homogenous but actually highly structured. We can place limits on the size of these spatial structures by considering the projection of the line of sight through the disk axis. Consecutive measurements at phases 0.964 and 0.975

yielded count rate variations of a factor of 8, these two measurements correspond to lines of sight passing through the disk axis $1.7 \times 10^9 m$ (or about 3% of the circumstellar disk radius) apart. We can therefore conclude that significant density fluctuations are present in the disk at radii of many R^* . In particular, we note that there are only large variability in flux/absorption when the NS moves away from us, not when it moves towards us. This raises the possibility that the NS may be producing the inhomogeneities in the circumstellar disk while it is intrinsically smoother beforehand. Similar effects are proposed for supergiant systems which also exhibit rapid changes in the absorption column before eclipses (Haberl, White & Kallman (1989) and Blondin et al. (1990)).

Finally, by comparing the long term decline in the X-ray flux with that of the optical, both after MJD 56580, it is possible to compare the two levels of flux change. From Equation 6 and Figure 14 it can be seen that the X-ray flux falls off at a significantly faster rate than the optical flux. In fact, assuming a baseline optical flux for no circumstellar disk of $I=16.1$, the optical falls by $\sim 13\%$ over 100d compared to $\sim 67\%$ for the X-ray. Since the X-ray flux is relying completely upon accreting material from the outer edges of the circumstellar disk, it is to be expected that a shrinking disk will, at some point, no longer be able to feed the accretion process whilst still retaining significant I-band emission. So a modest change in the disk size (and hence I-band emission) produces dramatic changes in the X-ray flux.

5 CONCLUSIONS

This paper reports the discovery of the first eclipsing Be/X-ray system in the Magellanic Clouds. Because of the eclipsing nature, an accurate orbital model solution has been possible. In addition, detailed observations have been obtained showing the neutron star being occulted by not only the companion Be star, but also some extended structure. We have interpreted this extended structure as the circumstellar disk around the Be star and have presented one possible simple model for how this might be configured. Clearly further exploration of all the orbital parameters combined with a physically more realistic model of the circumstellar disk is needed, but beyond the scope of this initial paper.

6 ACKNOWLEDGEMENTS

The OGLE project has received funding from the European Research Council under the European Community's Seventh Framework Programme (FP7/2007-2013)/ERC grant agreement no. 246678 to AU. VAM acknowledges funding from the NRF, South Africa. LJT and ESB are Claude Leon Foundation Research Fellows. Some of the observations reported in this paper were obtained with the Southern African Large Telescope (SALT).

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